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THE INFLUENCE OF SOLAR SPECTRAL LINES ON ELECTRON CONCENTRATION IN TERRESTRIAL IONOSPHERE

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Abstract. One of methods of detection and analyses of solar flares is by observing time variations of certain solar spectral lines. During solar flares, a raise of electron concentration occurs in Earth's ionosphere which results into amplitude and phase variations of recorded very low frequency (VLF) waves. We compared the data obtained by the analysis of recorded VLF signals and line spectra for different solar flares. In this paper we treated the DHO VLF signal transmitted from Germany at the frequency of 23.4 kHz recorded by the AWESOME system in Belgrade (Serbia) during solar flares in the period between 10:40 UT and 13:00 UT on April 22, 2011.

Key words: solar flares, line spectra, terrestrial ionosphere

1. INTRODUCTION

Solar radiation has a dominant external influence on the sunlit Earth's atmosphere. The rate of photo-ionization processes in the ionosphere depends on composite particles as well as on the solar radiation spectrum at the considered altitude. The data on some solar spectral lines could be of significant importance in studies of solar flares (see Valníček & Ranzinger (1972) and references therein). For example, Valníček & Ranzinger (1972) studied the observed widths of the H α line for different types of flares.

During the transmission through the terrestrial atmosphere, the solar radiation is being attenuated and, at the altitude of about 70 km, only X rays and Lyman α line remain important i.e. with sufficient energy to induce a noticeable photo-ionization of particles in the ionospheric D-region. The main subject of this paper is a study of the intensity of these lines and comparison of their time dependences in cases of unperturbed and perturbed ionosphere during different solar flares with the corresponding electron concentrations.

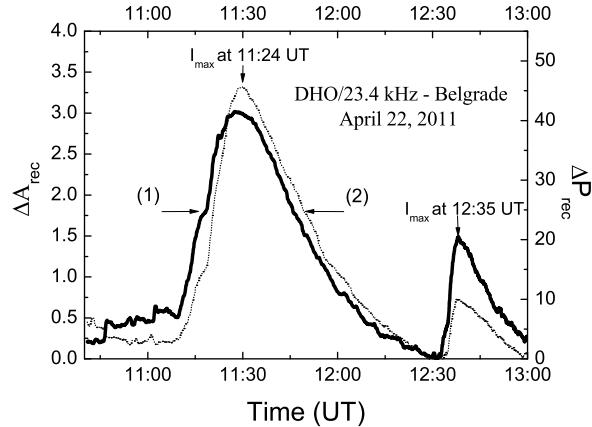


Fig. 1. Perturbed amplitude (1) and phase (2) of a signal emitted from DHO transmitter (Germany) and recorded by the AWESOME receiver in Belgrade (Serbia) during the observed flares. Null values correspond to the amplitude and phase recorded for the unperturbed ionosphere.

The electron concentration of plasma was calculated by using the developed and elaborated method in Grubor et al. (2008), Zigman et al. (2007) and Sulic et al. (2010). This calculations are based on electron concentration influence on propagation properties of very low frequency (VLF) waves. We used the time dependent solar spectrum during solar flares according to the model presented in Chamberlin et al. (2008).

2. THEORY AND RESULTS

Within the time interval 10:40 - 13:00 UT on April 22, 2011, the DHO signal (transmitted from Germany at frequency 23.4 kHz) recorded at Belgrade station has two dominant amplitude and phase increases. The measurable features of the VLF signal refers to the change in amplitude ΔA_{rec} , measured in dB, relative to the ambient levels prior to the event. The associated phase change ΔP_{rec} is measured too (Figure 1).

To calculate the electron concentration $N(t, h)$ in the D-region at altitude h , we used the equation (Wait 1964):

$$N(t, h) = 1.43 \cdot 10^{13} e^{-0.15H'(t)} e^{(\beta(t)-0.15)(h-H'(t))}. \quad (1)$$

The parameters of the Wait's model of ionosphere, the reflection height $H'(t)$ and sharpness $\beta(t)$, are calculated by the LWPC computer program (Ferguson 1998). The time distributions of the electron concentration at four altitudes, where the dominant electron gain and electron loss processes are the photo-ionization and recombination, respectively, are presented in Figure 2. One can see that at higher altitudes the photo-ionization processes strongly dominate the recombination ones. Enhancements of $N(t, h)$ result from increased solar radiation but the question remains which particular lines in solar irradiance spectrum are important for the photo-ionization processes in D-region. In introduction, we said that these can be the lines from the X spectrum and Lyman α line. The time distribution of these

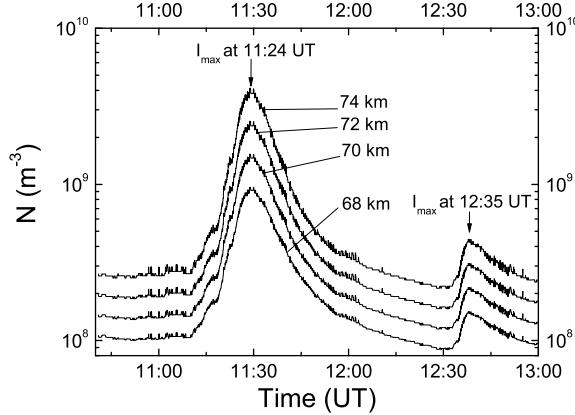


Fig. 2. Electron concentrations on different altitudes related to two flares of classes C7.7 and C3.0.

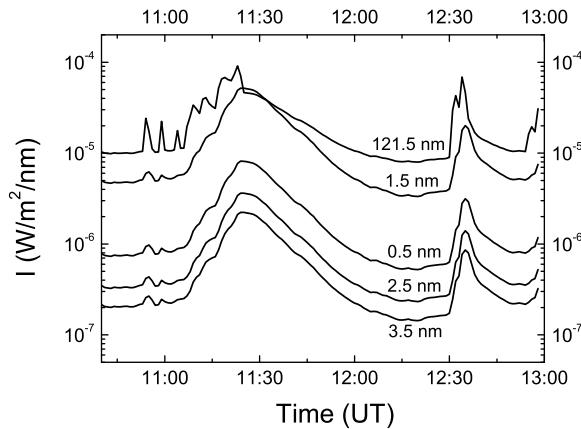


Fig. 3. Time distribution of the solar radiation intensity I which has the dominant role in perturbing the ionospheric D-region.

lines is presented in Figure 3. Here we can see that two X-flares and two increases of the Lyman α line intensity (121.5 nm curve) occurred within the observed time interval. Comparing Figure 2 with Figure 3 we can notice that the peaks in the time distribution of the electron concentration follow the time variation pattern of the X-emission spectrum but not that of the Lyman α line intensity. This leads to a conclusion that the influence of the X-spectrum lines is dominant for electron concentration at altitudes of about 70 km. Two intensity maxima of all considered X-spectrum lines occur at 11:24 UT and 12:35 UT, which is 5 min before the time the corresponding maxima appear in the electron concentration curve.

The spectrum of solar radiation below 150 nm, presented in Figure 4, shows that the analyzed lines in Figure 3 exhibit a bigger rise of intensity than the other

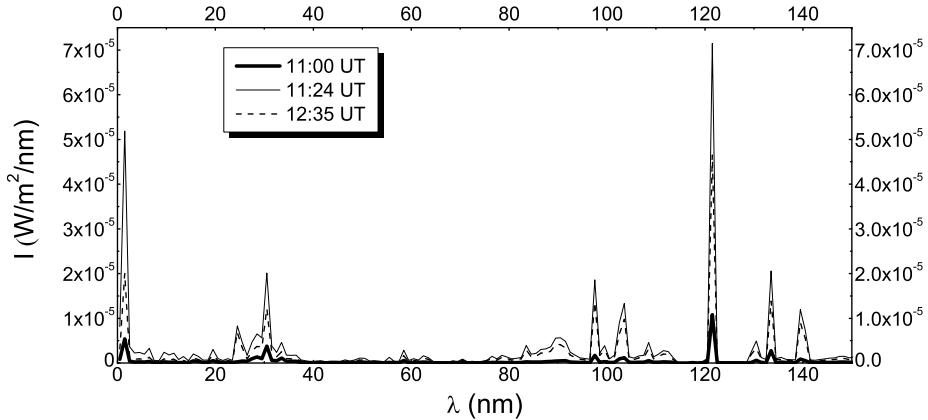


Fig. 4. Spectrum of solar radiation for the unperturbed ionosphere (dot line), at the maximum of the first (dash line), and at the maximum of the second (solid line) flare.

lines in the time interval of two maximal X ray intensities occurring at 11:24 UT and 12:35 UT. Also, we can see that the first X-flare produces higher line intensities which followed by a higher rise of electron concentration.

3. CONCLUSIONS

The analyzes presented in this work yield a conclusion that the dominant influence on enhancement of the electron concentration in the D-region during the observed solar flares is due to the increased intensity of radiation lines in the X-spectrum, i.e. X-flares with bigger line intensities produce larger rises of electron concentration. This conclusion follows from the fact that the rate of the photo-ionization process is higher then that of the recombination processes within the time interval between the start of the irradiance increase and to a few minutes after the irradiance maximum. In the case of these two flares, the maximum of electron concentration occurred after the appearance of the maximum intensity of lines produced by the X-flares.

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